

Magnetogenesis in bouncing cosmology postprint

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Abstract

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Full Text

Preamble

Magnetogenesis in Bouncing Cosmology

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energy scale of the bounce cosmology to the number of effective e-folding of the contracting phase in the case of scale invariance for the power spectrum of primordial magnetic fields. We show that this constraint can be satisfied in a sizable region of the parameter space for the nonsingular bounce cosmology.

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Introduction

Astronomical observations indicate the existence of magnetic fields throughout the observable universe—from the length scales of stars up to the scales of galactic clusters. The strength of these fields in galaxies and clusters is on the order of micro-Gauss. Magnetic fields can be produced from the amplifications of initial seeds via physical processes during structure formation (such as adiabatic compression and turbulent shock flows). Observations suggest the strengths of initial seed fields should be greater than 10^{-13} Gauss on a comoving scale larger than 1 Mpc [?, ?].

If the seed magnetic fields originated in the very early universe, before the surface of last scattering, they may affect the temperature and polarization spectra of the cosmic microwave background (CMB) anisotropies since their energy-momentum tensor can source the corresponding scalar, vector, and tensor cosmological perturbations. Therefore, CMB observations can provide experimental constraints on the properties of primordial magnetic fields [3–9]. According to the latest Planck 2015 data, primordial magnetic field strengths are required to be less than a few 10^{-9} gauss at the length scale of 1 Mpc [?].

Discovering the underlying mechanism responsible for generating primordial magnetic seed fields is an active area of cosmological research. In the literature, the origin of primordial magnetic fields has been widely discussed; see Refs. [11–13] for relevant reviews. Among these mechanisms, one possibility is that the seed fields could be produced during cosmic phase transitions such as the electroweak or QCD phase transitions. This approach, however, suffers from a manifest issue that the correlation lengths are too small when compared with observations. Consequently, one may investigate the possibility of producing seed fields at an even earlier time, such as during the inflationary epoch, in which case the process is known as inflationary magnetogenesis. Since the theory of electrodynamics is conformally invariant, and the energy density of the background universe is nearly constant, inflationary magnetogenesis requires the breaking of conformal invariance by introducing couplings of the electromagnetic fields to other sectors such as gravity [14–18], inflaton fields [19–24], pseudo scalar fields [14,25–27], or other fields [28–30]. However, these models typically produce seed fields that have too small strengths or suffer from other conceptual issues, such as the backreaction problem where the energy density of the generated electromagnetic fields spoils the inflationary background dynamics, and the strong coupling problem where the coupling of electromagnetism (EM) may become too large in the primordial stage. Addi-

tionally, there exists the curvature perturbation problem wherein the curvature perturbations seeded by the generated electromagnetic fields are too large to be in agreement with the CMB data \cite{31-36}. Recently, two of the current authors proposed a generalized model within the inflationary scenario in which the above issues were avoided under a certain fine tuning [?].

One possible alternative to the inflationary paradigm is that our observed universe may have evolved from a previously contracting phase and then experienced a nonsingular bouncing phase at very early times. This is known as a bouncing scenario \cite{38-43}. In particular, it has been shown that dynamical bounce models under certain contracting phases, such as the matter contraction \cite{44-46} or the ekpyrotic phase [?, ?, ?], can achieve a nearly scale invariant power spectrum of primordial fluctuations in order to explain the present cosmological observations. Various model constructions and related perturbation analyses were extensively performed in the literature, namely, the quintom bounce \cite{49-51}, the Lee-Wick bounce [?], the Horava-Lifshitz gravity bounce \cite{53-55}, the $f(T)$ teleparallel bounce \cite{56-58}, the ghost condensate bounce [?], the Galileon bounce [?, ?], the matter-ekpyrotic bounce \cite{62-64}, the fermionic bounce [?, ?], etc. (see, e.g., [?, ?] for recent reviews). In general, it was shown that on length scales larger than the time scale of the nonsingular bouncing phase, primordial cosmological perturbations remain almost unchanged through the bounce [?, ?]. The bouncing phase embedded in inflationary scenario was recently considered to explain both the large hemispherical power asymmetry and the power deficit on large angular scales in cosmic microwave background radiation [?, ?] (see [?, ?] for bouncing inflation).

In this paper we investigate the process of magnetogenesis within the framework of nonsingular bounce cosmology. Magnetogenesis in a bouncing universe was once investigated in [?, ?] where a specific form of the gauge field coupling was considered. Different from the case in inflationary scenario, the bounce cosmology suggests that our universe had a sufficiently long period of contraction and then experienced a smooth bouncing phase that connected the contracting phase with the observed expanding one. Correspondingly, the energy density of the background universe cannot be approximated to be a constant as in the inflationary scenario. It is notoriously difficult to construct bouncing models in General Relativity in a flat universe due to the required violation of the null energy condition (NEC). In general, NEC violation in bouncing cosmology is accompanied by a number of stability problems. If the matter Lagrangian is simply a general function of a scalar field and its derivative (as in k-essence) one cannot transition into a NEC violating phase and hence a bounce is not possible [?, ?]. A construction of a bounce which was manifestly free from gradient and ghost instabilities before thermal expansion and which occurs on trajectories of a non-vanishing measure was achieved in [?]. As a specific model for bouncing cosmology we choose [?] to study magnetogenesis; however, our findings are applicable to general bouncing models. In this paper we show that a bouncing universe can realize magnetogenesis in an efficient way due to the following three conditions. Firstly, as previously mentioned, there exists the

large scale problem for the generation of primordial magnetic field. This issue is nicely addressed within the bouncing scenario since in the contracting phase the wavelength of the seed fields exit the Hubble horizon. Secondly, the scale factor in the contraction phase is decreasing and the power spectrum of the magnetic field becomes dense instead of being diluted as in inflation. Thirdly, the energy density of the background universe also becomes dense in the contracting phase, and therefore the backreaction problem can be avoided within a reasonable tuning of model parameters.

The paper is organized as follows. In Section II we introduce the basic formulae of the model to realize magnetogenesis in bouncing cosmology. In Section III we briefly review the scenario of bounce cosmology by considering a specific model but keep the background evolution general. We then analyze the generation of primordial seed fields in Section IV and discuss constraints from the strong coupling issue. The backreaction issue of primordial magnetic fields is investigated in Section V within the bouncing scenario. We summarize our results in Section VI.

II. The Model of Bounce Magnetogenesis

To begin, we consider the following action for scalar field ϕ and electromagnetic field A_μ , minimally coupled to Einstein gravity:

$$S = \int d^4x \sqrt{-g} \left[\frac{1}{2} R + \mathcal{L}_{\text{bounce}} - \frac{1}{4} f^2(\phi) F_{\mu\nu} F^{\mu\nu} \right]$$

where $\mathcal{L}_{\text{bounce}}$ is the Lagrangian density responsible for the nonsingular bouncing background which will be introduced in detail in Section III, and the field strength $F_{\mu\nu}$ is defined in terms of the $U(1)$ gauge field in the familiar way $F_{\mu\nu} = \partial_\mu A_\nu - \partial_\nu A_\mu$. Unless explicitly stated, we will work with natural units in which $c = \hbar = M_{\text{pl}} = 1$. Moreover, the standard Lagrangian of the electromagnetic field can be recovered when the coefficient f is taken to be unity. In order to achieve the amplification of the magnetic fields, one expects the coefficient f to be a function of the background scalar field ϕ and hence can evolve along with the cosmic time so that the conformal symmetry is violated under this time evolution.

Varying the above action (1) with respect to the electromagnetic field yields the corresponding equation of motion:

$$\partial_\mu (\sqrt{-g} f^2 F^{\mu\nu}) = 0.$$

We consider an observationally favored, spatially flat Friedmann-Robertson-Walker metric,

$$ds^2 = a^2(\eta)(-d\eta^2 + d\mathbf{x}^2),$$

where $a(\eta)$ is the scale factor and η is the conformal time.

We choose the Coulomb gauge: $A_0 = 0$, $\partial_i A^i = 0$, and expand A_i in Fourier space as follows:

$$A_i(\eta, \mathbf{x}) = \int \frac{d^3 k}{(2\pi)^{3/2}} \sum_{\sigma=1,2} [\epsilon_{i,\sigma}(\mathbf{k}) a_{\mathbf{k},\sigma} A_k(\eta) e^{i\mathbf{k}\cdot\mathbf{x}} + \text{h.c.}],$$

where $\epsilon_{i,\sigma}(\mathbf{k})$ are two orthonormal and transverse polarization vectors. The operators $a_{\mathbf{k},\sigma}$ and $a_{\mathbf{k},\sigma}^\dagger$ are the annihilation and creation operators that satisfy the standard commutation relations $[a_{\mathbf{k},\sigma}, a_{\mathbf{k}',\sigma'}^\dagger] = \delta_{\sigma\sigma'} \delta(\mathbf{k} - \mathbf{k}')$ as well as $[a_{\mathbf{k},\sigma}, a_{\mathbf{k}',\sigma'}] = [a_{\mathbf{k},\sigma}^\dagger, a_{\mathbf{k}',\sigma'}^\dagger] = 0$.

Within the metric (3), the Fourier mode A_k with a fixed conformal wave number k satisfies the following equation of motion:

$$A_k'' + 2\frac{f'}{f} A_k' + k^2 A_k = 0,$$

with the normalization condition

$$A_k(\eta) A_k'^*(\eta) - A_k^*(\eta) A_k'(\eta) = \frac{i}{f^2},$$

being satisfied. In terms of a new variable $v_k = f A_k$, the equation of motion (5) can be reformulated as:

$$v_k'' + \left(k^2 - \frac{f''}{f} \right) v_k = 0.$$

From this equation, one can explicitly read that the evolution of the magnetic field, which is depicted by the gauge field A_i , depends on the form of the coupling coefficient f and the background evolution. In the following section we investigate the background of a bouncing universe.

III. Bounce Cosmology

Nonsingular bounce cosmologies appear in many theoretical frameworks where the gravitational sector is modified, by making use of Null Energy Condition violating matter fields, or in a universe with non-flat spatial geometry (see e.g. [?, ?]). However, it is important to note that the BKL instability could appear in the contracting phase since the back-reaction of anisotropies would dominate over the regular dust and radiation densities, unless one finely tunes the initial conditions to be nearly perfectly isotropic. This issue can be avoided if a period

of the ekpyrotic phase is introduced before the bounce [?] where the unwanted anisotropies are depressed in the presence of an ekpyrotic scalar field.

Recently, it was shown that nonsingular bouncing cosmology with regular matter contraction can be achieved with an era of ekpyrotic contraction by introducing a scalar field with a Horndeski-type, non-standard kinetic term and a negative exponential potential [?]. One could include a regular dust component and then assume the universe began in a state of matter-dominated contraction thus combining the matter bounce with the ekpyrotic scenario. The backreaction of anisotropies is manageable throughout the entire cosmological evolution of the matter-ekpyrotic bounce model, including at the bounce point [?], and hence, the BKL instability that arises for a large family of non-singular bounce models, as pointed out in [?, ?], is nicely resolved. A concrete realization of the matter-ekpyrotic bounce was constructed in [?] that involves two matter fields. This effective field theory model of a non-singular bounce can be embedded into a supersymmetric version [?], or into loop quantum cosmology [?] (also see [83–85] for different cosmological scenarios).

In the present study, we follow the original bounce model within the context of effective field approach that was proposed in [?] and its cosmological implications to magnetogenesis. The Lagrangian can be written as

$$\mathcal{L}_{\text{bounce}} = -K(\phi, X) - G(\phi, X)\square\phi,$$

where K and G are functions of the dimensionless scalar field ϕ and its canonical kinetic term $X = \partial_\mu\phi\partial^\mu\phi/2$. The d' Alembertian operator is given by $\square\phi \equiv g^{\mu\nu}\nabla_\mu\nabla_\nu\phi$.

Following [?] the function K is chosen as

$$K(\phi, X) = (1 - g(\phi))X + \beta X^2 - V(\phi),$$

and G as

$$G(X) = \gamma X,$$

where β, γ are the related model parameters. For a phase of ekpyrotic contraction, the form of the potential $V(\phi)$ is taken to be

$$V(\phi) = -V_0 e^{-\sqrt{2q}\phi} + e^{-b_V\sqrt{2q}\phi},$$

where V_0, q, b_V are model parameters as well. In order to obtain a violation of the Null Energy Condition near the end of the contraction phase, we take $g(\phi)$ to be

$$g(\phi) = g_0 e^{-\sqrt{2p}\phi} + e^{-b_g \sqrt{2p}\phi},$$

where g_0, b_g and p are model parameters whose reasonable range can be found in reference [?].

The evolution equations of the background dynamics are evaluated explicitly in this model, and are well studied in the literature. We do not repeat this study here, but rather provide a summary of the three stages of evolution most relevant to the process of magnetogenesis:

- the ekpyrotic contraction,
- the bouncing phase,
- the fast-roll phase.

Note that, compared to the first phase, the effects of the latter two are rather limited based on the analyses in [?], and hence, we treat them as instantaneous for simplicity. Moreover, we are interested in the contraction phase, during which the wavelength of the fields exits the Hubble horizon since this length scale contracts much faster than the scale factor. It is important to notice that during the same stage the energy density of the scalar field ϕ evolves as a perfect fluid with an almost constant equation of state w .

During this phase $\phi \gg -1$ so that $g \rightarrow 0$ and $\dot{\phi} \gg 1$. The Lagrangian for ϕ approaches the standard canonical form. There is an attractor solution for ϕ which yields an effective equation of state, $w \simeq -1 + \frac{2}{3q}$. If q has a critical value, the equation of state w can be sufficiently large and the energy density of ϕ will dominate over matter and radiation during contraction. The background in this case can be described as

$$a(\eta) \propto (\tilde{\eta}_e - \eta)^{\frac{q}{1-q}},$$

where the subscript e stands for the end of the contraction phase, $\tilde{\eta}_e = \eta_e - \eta$. η_e is the end of the contraction phase and $\eta = 0$ is the midpoint of the bounce scenario. The above description is valid until we reach η_e . For simplicity, we define $\tilde{q} = q/(1-q)$ and the scale factor can be expressed as

$$a(\eta) = a_e \left(\frac{\tilde{\eta}_e - \eta}{\tilde{\eta}_e - \eta_e} \right)^{\tilde{q}},$$

where a_e is the scale factor at the end of the contraction phase.

After the moment η_e , the universe can evolve through the nonsingular bouncing phase smoothly. During this phase, the evolution of the Hubble parameter H can be approximated as a linear function of the cosmic time t [?, ?]. Accordingly, the scale factor behaves as $a \sim \exp(t^2)$ roughly. In this case, one can read that a nonsingular bounce can occur when the scale factor reaches the minimal value.

In the following section we perform the analysis of the generation of the primordial magnetic fields in the aforementioned cosmological scenario.

IV. Bounce Magnetogenesis

After the bounce we wish to recover classical electromagnetism in order to connect with the observed big bang history smoothly, and therefore we require $f \rightarrow 1$ along the background evolution. If in the contraction phase f is much smaller than unity we will face the so-called strong coupling problem. A natural way to solve such a problem is to require $f^2 > 1$ during the contraction phase and it becomes unity at the end moment of the contracting phase η_e , which is also the beginning of the bouncing phase. This can be done by assuming that, during the contracting phase when $\eta < \eta_e$, f is a power-law function of the scale factor via

$$f = f_e \left(\frac{a}{a_e} \right)^n,$$

where $f_e \approx \mathcal{O}(1)$. Since in the contraction phase the scale factor decreases, we take $n > 0$ so that $f > 1$ is fulfilled.

During the contraction phase, the Hubble horizon decreases faster than the scale factor so that modes with larger wavelengths will exit the Hubble horizon earlier.

For short wavelength modes, or $k^2 \gg f''/f$, the solution to Eq. (7) is approximated as $v_k(\eta) = \frac{e^{-ik\eta}}{\sqrt{2k}}$, while long wavelength modes ($k^2 \ll f''/f$) have the general solution,

$$v_k(\eta) = C_1 f + C_2 f \int \frac{d\eta}{f^2},$$

where C_1 and C_2 are integration constants that can be fixed by matching this solution to (17). With the background $a \propto (\tilde{\eta}_e - \eta)^{\tilde{q}}$ and integrating the second term we find

$$v_k(\eta) = C_1 a^n + C_2 a^{-1-n+1/q}.$$

If $-1-n+1/q < n$, the second term will dominate over the first in the contraction phase. Using the conjunction condition $|v_k| \simeq \frac{e^{-ik\eta_k}}{\sqrt{2k}}$, where a_k stands for the scale factor when the corresponding wavelength exits the Hubble horizon $a_k |H| = k$ (which is also equivalent to $k(\tilde{\eta}_e - \eta_k) = \tilde{q}$), one derives

$$|C_2| \simeq \frac{1}{\sqrt{2k}} \left(\frac{a_k}{a_e} \right)^{n+1-1/q}.$$

We arrive at the complete form for v_k ,

$$v_k(\eta) \simeq \frac{1}{\sqrt{2k}} \left(\frac{a_k}{a_e}\right)^{n+1-1/q} \left(\frac{a}{a_e}\right)^{-n-1+1/q}.$$

Actually, with the scale factor given by Eq. (15) and the coupling function given by Eq. (16), the analytical solution to Eq. (7) is a linear combination of Bessel functions [?]. The two integration constants can be fixed by imposing the Bunch-Davies vacuum condition (17) as $(-k\eta) \rightarrow \infty$. For modes well outside the Hubble radius ($k^2 \ll f''/f$), one can get the same solution as (22) in the case of $-1 - n + 1/q < n$. Therefore, the gauge field is found to be

$$A_k = \frac{v_k}{f} = \frac{1}{\sqrt{2k}} \frac{1}{f_e} \left(\frac{a_k}{a_e}\right)^{n+1-1/q} \left(\frac{a}{a_e}\right)^{-2n-1+1/q}.$$

Correspondingly, the power spectrum of the primordial magnetic fields is evaluated as

$$P_B(k, \eta) = \sum_{\sigma=1,2} \frac{k^5 |A_k(\eta)|^2}{2\pi^2 a^4} = \frac{k^4}{\pi^2 f_e^2} \left(\frac{a_k}{a_e}\right)^{2\tilde{q}(n+1-1/q)} \left(\frac{a}{a_e}\right)^{-4n-6+2/q},$$

with the spectral index $n_B = (6 - 2(3 + n)q)/(1 - q)$. Note that the spectrum is scale-invariant for $n = -3 + 3/q$. At the end of the contraction phase we have

$$P_B(k, a_e) = \frac{k^4}{\pi^2 f_e^2} \left(\frac{a_k}{a_e}\right)^{2\tilde{q}(n+1-1/q)}.$$

As argued previously, the effects of the (fast) bounce are limited and hence we treat this phase instantaneously. Note that the vector field has become the standard electromagnetic field right before the bounce at the moment η_e . After that, the magnetic fields decay with the square of the scale factor. The strength of the magnetic field observed today is given by:

$$B_0 = \sqrt{P_B(k, a_0)} = \frac{k^2}{\sqrt{\pi} f_e} \left(\frac{a_k}{a_e}\right)^{\tilde{q}(n+1-1/q)} \left(\frac{a_0}{a_e}\right)^{-2},$$

where a_0 is the scale factor today. In the scale-invariant case we have

$$B_0 \simeq \frac{H_e^2}{\sqrt{\pi} f_e} \left(\frac{a_e}{a_0}\right)^2.$$

We must determine how much the scale factor changes from the end of the contraction until today. Using the conservation of entropy we find

$$\frac{a_e}{a_0} \sim \frac{T_0}{H_e} \exp(N),$$

where T_0 is the CMB temperature today and N is the effective e-folding number of the contraction phase.

We treat the energy density of the electric field as the energy density of radiation at the end of the contraction phase because it dominates over that of magnetic fields. The power spectrum of the electric fields takes the form of

$$P_E(k, \eta) = \sum_{\sigma=1,2} \frac{k^3 |A'_k(\eta)|^2}{2\pi^2 a^4} = \frac{(2n+1-1/q)^2 H^4}{2\pi^2 f_e^2} \left(\frac{a_k}{a_e}\right)^{2\tilde{q}(n+1-1/q)-2} \left(\frac{a}{a_e}\right)^{-4n-4+2/q}.$$

The ratio of P_E to P_B can be written as

$$\frac{P_E}{P_B} = (2n+1-1/q)^2 \left(\frac{a}{a_e}\right)^{2-2/q} \left(\frac{a_e H_e}{k}\right)^4,$$

from which we can see that $P_E \gg P_B$ for the long wavelengths beyond the Hubble horizon. Therefore, we can compute the energy density of the electric part as the total energy density of the electromagnetic fields:

$$\rho_{\text{EM}} = \frac{f^2}{2} \langle F_{\mu\nu} F^{\mu\nu} \rangle \simeq \int_{a_i H_e}^{a_e H_e} \frac{dk}{k} P_E(k, a_e).$$

This yields

$$\rho_{\text{EM}} = \frac{(-1+q+2nq)^2(1-q)(a_e H_e)^{2\tilde{q}(n+1-1/q)-2} H_e^4}{2\pi^2 q^2(4-2(2+n)q)} \left[\left(\frac{a_e H_e}{a_i H_i}\right)^{4-(4+2n)q/(1-q)} - 1 \right],$$

where a_i is the scale factor at the moment when the largest scale observed today first exited the Hubble horizon.

Since at the end of the contraction phase $a_i H_i \ll a_e H_e$ in the case of the scale-invariant power spectrum, the first term on the right-hand side of equation (31) can be ignored. Thus, we have

$$\rho_{\text{EM}}(a_e) \simeq \frac{25(1-q)^2 a_e^2 H_e^6}{2\pi^2 q^2 (a_i H_i)^2}.$$

From equations (32) and (28), we get

$$B_0 \simeq \frac{5f_e(1-q)}{\sqrt{2\pi q}} \frac{H_e^2}{(a_i H_i)} \left(\frac{a_e}{a_0}\right)^2 \simeq \frac{5f_e(1-q)}{\sqrt{2\pi q}} T_0 H_e \exp(-N),$$

where N is the effective e-folding number of the contraction phase. Following Ref. [?], it is limited by

$$N < \frac{1-q}{2(1-3q)} \ln \left(\frac{1}{\Omega_{\text{pert}}} \right),$$

where Ω_{pert} is the amplitude of the observed anisotropy in CMB. This leads to

$$B_0 < \frac{5f_e(1-q)}{\sqrt{2\pi q}} T_0 H_e \exp \left[-\frac{1-q}{2(1-3q)} \ln \left(\frac{1}{\Omega_{\text{pert}}} \right) \right].$$

Choosing $f_e = 1$ and $q = 0.1$, and using $T_0 = 1.9 \times 10^{-32}$ in Planck units and $\Omega_{\text{pert}} = 10^{-5}$, we have the upper bound of the strength of the magnetic fields today:

$$B_0 \lesssim 10^{-11} \text{Gauss.}$$

Note that we need to ensure that the contracting phase proceeds for a sufficiently long period as discussed above, and the condition that $-1 - n + 1/q < n$ has to be satisfied in the scale-invariant case. The combination of these two conditions then yields the upper bound of B_0 as a function of the model parameter q . In Fig. 1 [Figure 1: see original paper], we demonstrate the relation between the upper bound of B_0 and the parameter q for specific parameter choices.

V. Constraints from the Backreaction

In the above section we performed the analyses of magnetogenesis within bounce cosmology and showed that the corresponding results can satisfy the observational constraints which gives rise to an upper bound for the produced magnetic fields. In this section we further check the conditions that the energy density of the electromagnetic fields does not spoil the evolution of the background dynamics, which requires

$$\rho_{\text{EM}} < \rho_\phi,$$

where ρ_ϕ stands for the energy density of the background field. Since ϕ dominates over other components at late times, we may compare it with the contribution of the electromagnetic fields directly for simplicity. From (31) we find $\rho_{\text{EM}} \propto a^{-2n-4} \propto a^{2-6/q}$ in the case of scale invariance. Moreover, the energy density of the background scalar field evolves as $\rho_\phi \propto a^{-3(w+1)} \propto a^{-2/q}$.

The above two equations imply that the energy density of EM grows even faster than that of ϕ . We require that at the end of the contraction phase ϕ dominates, that is to say,

$$\rho_{\text{EM}}(a_e) < H_e^2 M_{\text{pl}}^2,$$

where we recover the Planck energy scale unit M_{pl} to balance the units on both sides of the above equation. This gives

$$\frac{25(1-q)^2 a_e^2 H_e^6}{2\pi^2 q^2 (a_i H_i)^2 M_{\text{pl}}^2} < H_e^2,$$

which leads to

$$\frac{|H_e|}{M_{\text{pl}}} < \frac{\sqrt{2}\pi q}{5(1-q)} \left(\frac{a_i H_i}{a_e H_e} \right) \simeq \frac{\sqrt{2}\pi q}{5(1-q)} \exp(-N).$$

As a result, the inequality yields a theoretical constraint on the relation of the effective e-foldings N to the energy scale of ϕ at the end of contraction phase.

In order to have a clearer picture of this theoretical constraint, we numerically depict the parameter space that is allowed for a well-behaved bounce model in Fig. 2 [Figure 2: see original paper]. We find this theoretical constraint is easily satisfied in a sizable region of the parameter space of nonsingular bounce cosmologies.

VI. Conclusion

In this paper we showed it is possible to produce appreciable primordial magnetic fields within the context of nonsingular bounce cosmology. We adopted the specific model developed in [?] but the results derived in our study are applicable to more general bouncing models as the main analyses only depend on the phase of ekpyrotic contraction. In this model, we found that primordial magnetic field with potentially sizable strength could be produced without troublesome strong coupling and backreaction effects, which plague the process of magnetogenesis in other very early universe paradigms such as inflation. By focusing our interest on the nearly scale-invariant power spectrum of primordial magnetic fields, we derived an explicit observational constraint upon the relation between the magnetic field that may be observed today and the model parameter that is associated with the effective equation of state during the contracting phase. Moreover, by examining the limit of the backreaction, we obtained a theoretical constraint on the relation of the energy scale of the bounce cosmology with respect to the effective e-folding numbers of the contracting phase in the case of scale invariance. We showed that this theoretical constraint can be satisfied in a sizable parameter space of bounce cosmologies.

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